



# Correction of errors in scale values for magnetic elements for Helsinki

L. Svalgaard

Stanford University, Cypress Hall C3, 466 Via Ortega, Stanford, CA 94305, USA

Correspondence to: L. Svalgaard (leif@leif.org)

Received: 2 November 2013 – Revised: 2 April 2014 – Accepted: 28 April 2014 – Published: 5 June 2014

**Abstract.** Using several lines of evidence we show that the scale values of the geomagnetic variometers operating in Helsinki in the 19th century were not constant throughout the years of operation 1844–1897. Specifically, the adopted scale value of the horizontal force variometer appears to be too low by  $\sim 30\%$  during the years 1866–1874.5 and the adopted scale value of the declination variometer appears to be too low by a factor of  $\sim 2$  during the interval 1885.8–1887.5. Reconstructing the heliospheric magnetic field strength from geomagnetic data has reached a stage where a reliable reconstruction is possible using even just a single geomagnetic data set of hourly or daily values. Before such reconstructions can be accepted as reliable, the underlying data must be calibrated correctly. It is thus mandatory that the Helsinki data be corrected. Such correction has been satisfactorily carried out and the HMF strength is now well constrained back to 1845.

**Keywords.** Geomagnetism and paleomagnetism (time variations, diurnal to secular) – interplanetary physics (interplanetary magnetic fields; instruments and techniques)

## 1 Introduction and rationale

After more than a decade of vigorous research (e.g., Lockwood et al., 1999; Svalgaard et al., 2003; Svalgaard and Cliver, 2005, 2010; Lockwood and Owens, 2011) the magnitude of the heliospheric magnetic field (HMF) near Earth is well constrained from 1883 (probably even from 1872) to the present during which period sufficient and accurate geomagnetic data is available for calculation of the IDV index (Svalgaard and Cliver, 2005, 2010) that serves as a proxy for the HMF strength,  $B$ . There is still a healthy debate about the reconstruction before 1883 when geomagnetic data becomes

sparse and subject to errors, especially in the more difficult to measure horizontal force.

Although the IDV-index is calculated from the unsigned difference between the horizontal force at consecutive local midnight hours, it was already pointed out by Svalgaard and Cliver (2003, 2005) and Svalgaard and Cliver (2010, Fig. 6) that the IDV index can be computed for any hour and for any geomagnetic element. Conforming with that stipulation, Lockwood and colleagues (e.g., Lockwood et al., 2013a, b; Lockwood, 2013) have suggested to reduce the influence of noise in the early 19th century geomagnetic data by computing the average of the 24 individual time series of IDV calculated for each of the 24 h of the day, dubbed IDV(1d). Although this procedure introduces unwanted variance because of the day-to-day variability of the (semiregular) diurnal variation of the geomagnetic field, the “IDV signature” is robust enough such as to reduce this extra variance to a second-order effect.

Nevanlinna and colleagues (Nevanlinna et al., 1992; Nevanlinna, 2004) have compiled archived geomagnetic observations from the Helsinki (IAGA – International Association of Geomagnetism and Aeronomy – designation HLS) magnetic observatory comprising over 2 million observations of  $H$  and  $D$  components measured during the interval 1844–1912 with time resolution of 10 min to 1 h. Because of disturbances from nearby electric tram lines and general curtailment of the observational program, reliable and complete daily records of hourly values are only available up through 1897. Lockwood et al. (2013a, b) used this HLS data set to calculate IDV(1d).

The validity of the resulting IDV(1d) series and of conclusions drawn from it, obviously hinges on the data being correctly calibrated. In this note we shall show that the scale value for the horizontal force is seriously incorrect (too

small) during the interval 1866.0–1874.5 and that the scale value for the declination is also incorrect (too small) during the interval 1885.8–1887.5. These errors must be corrected before further inferences are drawn from the Helsinki data and before the results can be compared with and integrated into other reconstructions, in order to extend to earlier times the undisputed consensus reached for HMF *B* after 1882.

## 2 The summed ranges

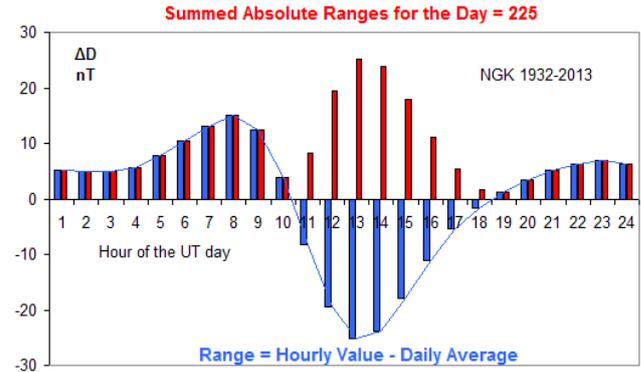
Bartels (1925, 1932) defined the *u* measure as the average (over from 1 to 12 months) unsigned difference between the daily means of the horizontal force, formally equivalent to the proposed IDV(1d) index and used by Svalgaard and Cliver (2005) in the derivation of their IDV index. Before 1872 there were no readily available daily mean data for any magnetic observatory, so Bartels – “more for illustration than for actual use” – turned to use the “summed ranges” (designated *s*) supplied by Moos (1910) as the main contributor to a proxy for the *u* measure.

Bartels’ interpretation of Moos’ procedure and data (Moos, 1910, Table 261) was “*s* is derived from the mean diurnal variation of *H* at Bombay for each single month, expressed in departures from the average, and is the sum of these departures, summed without regard to sign”. Using monthly means attenuates the irregular strong disturbances associated with the IDV signature and gives undue weight to the regular daily variation, thus downplaying the role of true “disturbances”. Moos was aware of this and on his effort of making a list of days classified as quiet or disturbed (ibid. page 421) remarked that “[for] a list of the kind [...] involving a large personal equation, some additional data are clearly essential in order to make the classification more mathematically definite. The daily range, or preferably the summed ranges, figures of the diurnal inequality of *each day*<sup>1</sup> would probably serve as the most appropriate data for this purpose; but as this is not possible on account of the heavy labour involved in their derivation”; so he resorted to the monthly means eventually used by Bartels. Here we shall build on that intuition (based on Moos’ extensive knowledge of the phenomenon) as we are no longer limited by computational power, provided data in digital tabular format is available. To make things explicit, Fig. 1 illustrates our interpretation of Moos’ prescription, emphasizing that by *s* we shall henceforth in this paper mean *s* derived from daily departures.

### 2.1 Calculating IDV from summed ranges

We begin by calculating *s* for both the declination, *s*(*D*), and for the horizontal force, *s*(*H*), for the German station Potsdam (POT, 1890–1907) and its replacement stations Seddin (SED, 1908–1931) and Niemegek (NGK, 1932–2012).

<sup>1</sup>Italics added.



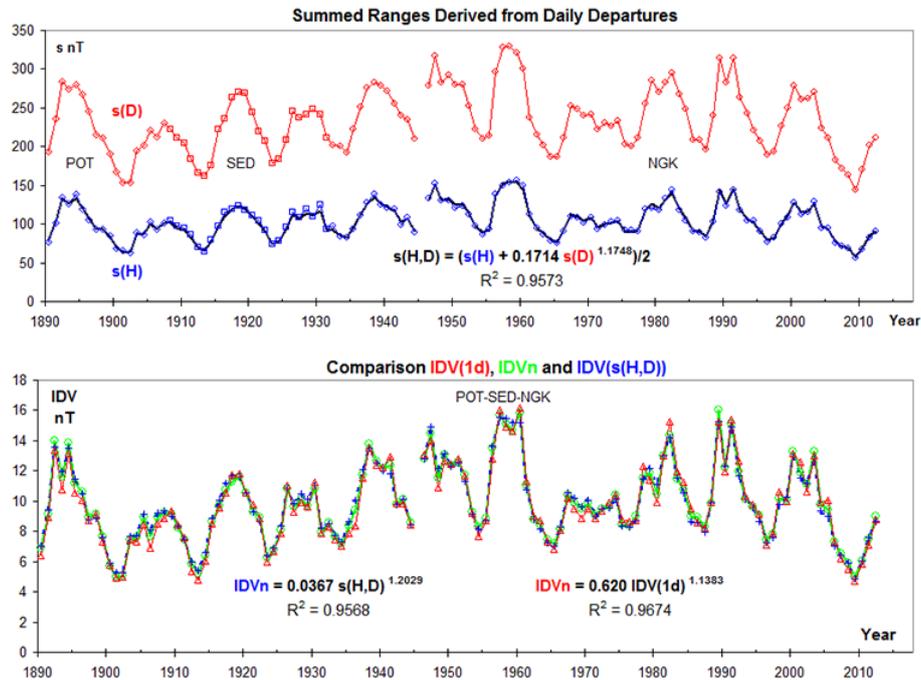
**Figure 1.** Average diurnal variation of declination (expressed in force units, nT) at Niemegek. On any given day, the variation consists of a pattern as shown here (although varying a bit from day to day) with superposed “noise” from geomagnetic activity, thus increasing the variance; this increase is what we are interested in. The signed deviations (blue bars determined every hour – either from an instantaneous value on the hour or from the hourly mean) from the daily mean are converted to unsigned departures (red bars), which are then summed over the day giving (as Moos expressed it) the summed ranges for each day, denoted by *s*.

Geomagnetic conditions were essentially the same at all three stations, because they were carefully placed with that in mind, so we can treat the data as homogeneous from a single station (Fig. 2).

By inspection it is clear that *s*(*D*) and *s*(*H*) are highly correlated, in fact  $s(H) = 0.1714s(D)^{1.1738}$  (the coefficient of determination  $R^2 = 0.9573$  is calculated from the linear fit of the logarithms); see also Fig. A1 in the Appendix. We can then form the average *s*(*H*, *D*) of observed *s*(*H*) and *s*(*H*) calculated from *s*(*D*) as shown by the black line in the upper panel of Fig. 2. We can calculate the IDV index for this particular station chain in the usual way using unsigned differences between the hour following local solar midnight, call it IDV<sub>n</sub>. The series of IDV<sub>n</sub> and *s*(*H*, *D*) are also highly correlated:  $IDV_n = 0.0367 s(H, D)^{1.2029}$  ( $R^2 = 0.9568$ ); Fig. A2 in the Appendix. It is rare to find correlations that significant. The conclusion is that given either *s*(*D*) or *s*(*H*) or both, we can calculate a very close approximation (blue curve) to the usual IDV<sub>n</sub> (green curve) as shown in the lower panel of Fig. 2. This is particularly important for early stations where *H* is often very noisy, while *D* is well-observed (or at times even the only component observed).

### 2.2 Calculating IDV from IDV(1d) or *u* measure

Because the *u* measure was found to be a good proxy for IDV<sub>n</sub> (e.g., Svalgaard and Cliver, 2005) we expect the equivalent IDV(1d) index (Lockwood et al., 2013a) to be as well. The red line in the lower panel of Fig. 2 bears this out:  $IDV_n = 0.620 IDV(1d)^{1.1383}$  ( $R^2 = 0.9674$ ) for this homogeneous data set, consistent with the finding by Mayaud (1980,



**Figure 2.** (Upper panel) Summed ranges derived from daily departures for declination  $s(D)$  (red curve) and horizontal force  $s(H)$  (blue curve) for the combined POT-SED-NGK series. Each station's yearly value is marked with a different symbol (POT diamond, SED square, NGK circle). The break in 1945 was caused by interruptions stemming from the Battle for Berlin during the final phase of WWII. The composite  $s(H, D)$  is added over  $s(H)$  as a black line. It is difficult to distinguish between the blue and the black lines. It is rare in this business to find such close agreement.

(Lower panel) IDVn (from midnight values) for POT-SED-NGK (green line) compared to IDV computed from  $s(H, D)$  (blue dashed line). Because the two curves are so close to at times be indistinguishable, each yearly value is also marked with a symbol: green circle for IDVn and blue plus sign for IDV ( $s(H, D)$ ). Finally, the red curve and red triangles show IDV(1d) scaled to IDVn as indicated. We need that scaling because IDV(1d) is about 11 % higher than IDVn due to the day-to-day variability of the regular daily variation.

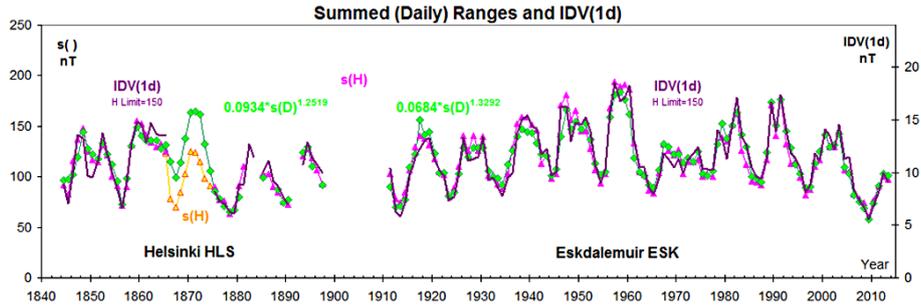
p. 13) that “the  $u$  index . . . certainly suffers from intrinsic defects . . . One might suspect a contamination by the regular variation, since its day-to-day variability should contribute to the interdiurnal variability. However, we tried to evaluate the importance of this contamination and were astonished at its relative smallness.” So, we have essentially three different ways of estimating IDV. This also holds for other long-term homogeneous station sets, e.g., PSM(Parc Saint-Maur)-VLJ(Val Joyeux)-CLF(Chambon-la-Forêt). As long as we limit ourselves to stations far enough ( $> 10^\circ$ ) from the auroral zone these three different methods yield comparable and highly correlated values.

### 3 Scale values for Helsinki data

The original archived data for Helsinki Observatory (situated at  $60^\circ 10.4' N$ ,  $24^\circ 56.9' E$ ) was given in “scale units”, which must be converted to force units (nT, called gammas in older literature) or angles (typically tenths of arc minutes). The scale units must be converted into physical units. The usual scheme calculates the physical values from the scale units like this: physical value = base value + scale value · (scale

units + instrument corrections). Often the base value and the instrument corrections are not known and the magnetometer can be characterized only as a “variometer”. The scale value must be known, either from instrument characteristics or from comparison with other instruments or other data, for the data to be of use.

The first director of the observatory, Nervander (1850), gave the scale value for the  $H$  variometer as  $3.6 \text{ nT mm}^{-1}$ . The  $H$  variometer is sensitive to temperature changes and the temperature was recorded, but the temperature coefficient was only determined later and the data then corrected for temperature variations. The scale value for the  $D$  variometer calculated from the characteristics of the instrument is quoted as  $0.315 \text{ arcmin mm}^{-1}$ . No absolute measurements are given and the scale values are assumed to be constant, as there is no known metadata about changes of calibration, which, however, does not necessarily mean that there were no such changes. We shall show here that significant changes took place and that the data need to be corrected accordingly, preferably from (as yet) undiscovered metadata, if possible, and if not possible, from comparison with other observatories and other solar or geomagnetic proxies, e.g., as in Morozova et al. (2014).



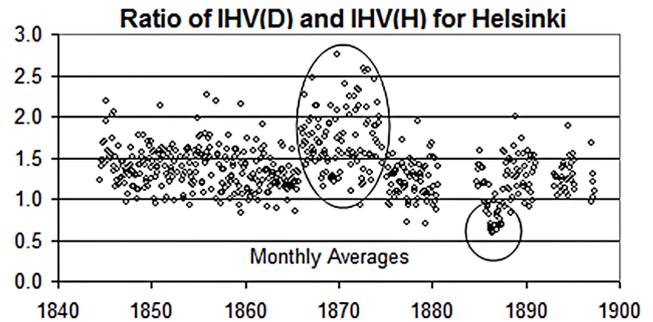
**Figure 3.** Yearly average summed ranges for  $H$  (pink triangles) and for  $D$  (green diamonds) scaled to match the scale of  $H$  using the equations in green for HLS (left) and ESK (right). The equations are slightly different because the inhomogeneous “raw” values are plotted, i.e., not normalized to a common “bridge”. The values of  $s(H)$  for the interval 1866–1874.5 (orange symbols) do not match the rest of the  $s(H)$  to scaled  $s(D)$ . IDV(1d) calculated from  $H$  for ESK (purple symbols; scale on right) is a good fit to  $s(H)$  and  $s(D)$ . A few “spikes” have been suppressed by capping daily IDV(1d) at 150 nT. IDV(1d) calculated from  $H$  for HLS (scale at right; same scale as for ESK as no normalization between HLS and ESK is performed – this is raw data – for this inhomogeneous data set) is also a good fit to  $s(H)$  and  $s(D)$ , except for the interval 1866–1874.5.

**3.1 Calculating IDV(1d) from  $s$  at ESK and HLS**

In Lockwood et al. (2013a, b) the IDV(1d) series for HLS and ESK (Eskdalemuir) are spliced together using POT as a “bridge”. In spite of the bridge being at considerably lower corrected geomagnetic latitude (by  $6^\circ$ ), it is posited that the result is a homogeneous data set. If so, results from ESK should be applicable to HLS as well. In Fig. 3 we show in the right-hand part the very similar variations since 1911 of the summed ranges  $s(H)$  and of  $s(D)$  (scaled to  $s(H)$ ) and of IDV(1d) for ESK. This is as expected from the results demonstrated in Sect. 2. For HLS, shown in the left-hand part of the figure,  $s(H)$  and (scaled)  $s(D)$  also agree closely, *except for the interval 1866–1874.5* (orange data points). The simplest, and in our view inescapable, conclusion to draw from this discrepancy is that the adopted scale value of the  $H$  variometer was too small, by some 30%, during the interval 1866–1874.5. Figure A3 in the Appendix compares IDV(1d) for HLS calculated from the summed ranges, further visualizing the obvious discrepancy. IDV(1d) for HLS, calculated for years outside of the interval 1866–1874.5 is plotted in Fig. 3 as well, for comparison. The agreement with  $s(H)$  and  $s(D)$  is again good.

**4 IHV also shows the scale value discrepancies**

Svalgaard and colleagues (Svalgaard et al., 2003, 2004; Svalgaard and Cliver, 2007) introduced the interhourly variability index, IHV, as a proxy for auroral zone activity (as measured at midlatitudes). Although HLS is too close to the auroral zone for IHV calculated from HLS data to retain its simple physical meaning (a proxy for solar wind  $BV^2$  and for the NOAA/POES hemispheric power; Emery et al., 2008), the IHV values do depend directly on the scale value used for the variometers. As for IDV, IHV can be computed for any geomagnetic element. If the scale values for  $H$  and for  $D$  were

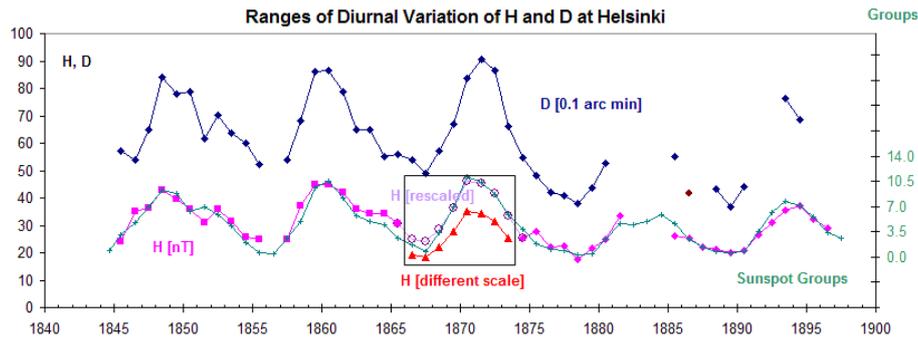


**Figure 4.** The ratio between monthly values of IHV calculated using the declination, IHV ( $D$ ), and of IHV calculated using the horizontal force, IHV ( $H$ ) for Helsinki. The ovals show the effect of the scale value for  $H$  being too low in the interval 1866–1874.5 and of the scale value for  $D$  being too low in the interval 1886–1887.5.

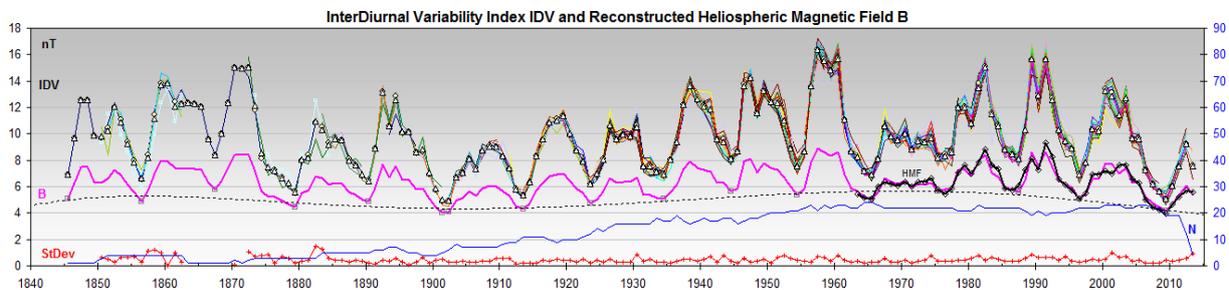
both correct, the ratio between IHV ( $D$ ) and IHV ( $H$ ) would be constant (apart from a random noise component). Figure 4 shows the ratio between IHV ( $D$ ) and IHV ( $H$ ) for HLS. As predicted from the analysis in Sect. 3 the ratio shows the expected behavior (in large oval) for the interval 1866–1874.5. The smaller oval shows that there is also a problem with the scale value of the  $D$  variometer in the interval 1886–1887.5, being too low by about a factor of 2. This is explored further in the Appendix (Fig. A5).

**5 The daily variation**

The diurnal variation ( $S_R$  sometimes less accurately called  $S_q$ ) of geomagnetic elements can be used to check the scale value of the magnetometers. Computing for each day the differences between the instantaneous hourly values (or hourly means – the distortion caused by averaging over an hour is but slight) and the daily mean removes the effect of the



**Figure 5.** Yearly average ranges for declination  $D$  (in 0.1 arc minute units), blue curve, and for horizontal force (in nT units), pink curve. Because of the strong seasonal variation only years with no more than a third of the data missing are plotted. The green curve (with “+” symbols) shows the number of active regions (sunspot groups) on the disk scaled to match the pink curves ( $H$ ). As expected the match is excellent, except for the interval 1866–1874, where the  $H$  range would have to be multiplied by 1.32 for a match, purple open circles.



**Figure 6.** Reconstruction of annual means of 169 years of near-Earth heliospheric magnetic field strength  $B$  (pink line in middle of graph) for the interval 1845–2013 compared with in situ spacecraft measurements (black line marked HMF) plotted using different colors for each station, from Svalgaard and Cliver (2014). Open triangles (or circles) show the median (or mean) of all stations in each year. The red line at the bottom of the graph shows the standard deviation of the values of IDV in each year. The blue line marked  $N$  shows the number of stations for each year.

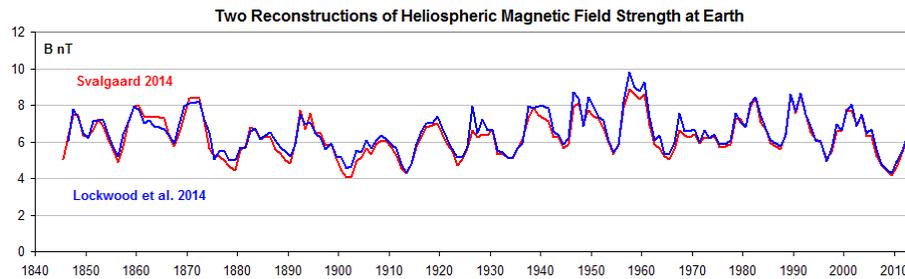
(slowly varying) secular values and of random (unknown) changes in the baseline. The average, over an interval – such as a month or a year – of the differences as a function of time within the day is the average diurnal variation (what used to be called the daily “inequality”). It is well-known that that average range, i.e., the difference between the maximum and minimum values of the average diurnal variation, is extremely well correlated with appropriate solar activity indices (e.g., F10.7 microwave flux, sunspot number, or the group number (number of active regions on the solar disk)), as was discovered by Wolf (1852)<sup>2</sup> and subsequently extensively verified by many workers (e.g., Bartels, 1946), considered to be the best of all solar–terrestrial correlations; a fact used by Nevanlinna et al. (1992), who note “(t)he scale value for the  $D$  variometer seems to be reliable (for 1844–1853) because the diurnal variation at the Helsinki, Nurmijärvi and St. Petersburg observatories show very similar

<sup>2</sup>“Wer hätte noch vor wenigen Jahren an die Möglichkeit gedacht, aus den Sonnenfleckenbeobachtungen ein terrestrisches Phänomen zu berechnen?” (Who would have thought, just a few years ago, of the possibility of computing a terrestrial phenomenon from observations of sunspots?)

behavior being the same within 1’ under corresponding solar activity conditions”. Figure 5 shows the yearly average ranges for  $D$  and  $H$  at Helsinki.

The group numbers used in Fig. 5 are derived from the recent re-evaluation of solar activity (Svalgaard, 2013). Using the official SIDC (Solar Influences Data Center) sunspot number does not change the result for the interval of interest (Appendix, Figs. A4, A5). It seems that the scale value adopted for  $H$  during the interval 1866–1874 must actually be different from that used for the rest of the  $H$  data, specifically 1.32 times lower than the constant value used by Nevanlinna (2004) in constructing the Helsinki series. The range of the declination during that interval matches that of  $H$  when  $H$  is rescaled upwards by the factor 1.32. The ranges of  $D$  and  $H$  generally vary together (with solar activity) due to the same current system, so the discrepancy indicates a problem with the adopted scale value of  $H$ . Figure A4 in the Appendix documents the problem for the  $H$  component for the year 1869, while Fig. A5 in the Appendix documents the problem for the declination for the year 1886.

An equally strong case can also be made comparing the diurnal range of  $H$  at Helsinki directly to that of  $H$  at other



**Figure 7.** Comparison of the HMF strength deduced by Svalgaard and Cliver (2014) as shown in Fig. 6 (red curve) and inferred by Lockwood et al. (2014a) (blue curve). The coefficient of determination is  $R^2 = 0.93$ .

stations, not using the solar activity connection. In Fig. A6 of the Appendix, we show a comparison with Greenwich (GRW, brown), Prague (PRA, blue), and Colaba (CLA and replacement station Alibag ABG, green). Because not all stations observed hourly values all the time, the ranges have been matched to Helsinki (HLS, pink) outside the interval 1866–1874. During that interval, the range for HLS (red triangles) is seriously too low.

## 6 Reconstructed HMF

Figure 6 shows a reconstruction of annual means of 169 years of near-Earth heliospheric magnetic field strength  $B$  (pink line in middle of graph) for the interval 1845–2013 compared with in situ spacecraft measurements (black line marked HMF). The reconstruction is based on a re-evaluation of the IDV index using the normalized average of the three determinations discussed in Sects. 2 and 3, from Svalgaard and Cliver (2014), plotted using different colors for each station. For the interval 1863–1871 leading up to the strong solar cycle 11, only HLS contributes (awaiting digitization of other stations), underscoring the importance of getting HLS right.

A consequence of the undue weight given to the regular diurnal variation that we referred to in Sect. 2 when using the summed ranges based on monthly averages that Bartels used to extend the  $u$  measure before 1872 is that our earlier reconstruction based on the  $u$  measure (Svalgaard and Cliver, 2010) for years with large coronal holes and ensuing large HMF during the declining phase of the solar cycles was too low for years during the declining phase. Going to the summed *daily* ranges,  $s(H, D)$ , remedies that defect, especially when the erroneous scale values are corrected. It is instructive to compare the reconstruction for solar cycles 10 and 11 (1857–1878) with those of cycles 18 and 19 (1945–1964), as regard to both the “shape” of the solar cycle curves and to the similar general level of activity.

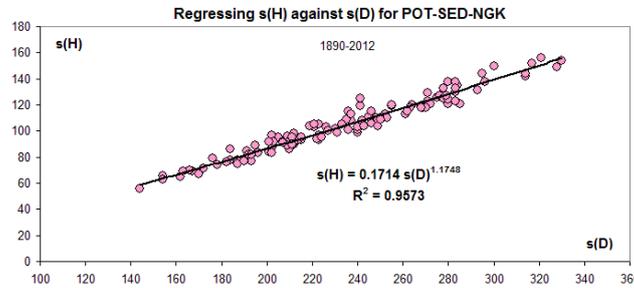
After the present paper was submitted, Lockwood et al. (2014a; 2014b) and Lockwood and Owens (2014), now aware of our finding, have accepted our analysis and corrected their reconstruction accordingly. Figure 7 shows that their revised values (Lockwood et al., 2014a) are largely correct (compared with our multistation reconstruction), and that reconstruction of the HMF strength is now satisfactory constrained back to 1845. A significant insight that follows from the concordant reconstructions is that there hardly was any “modern grand maximum” as the values of the HMF in the 20th century are on par with the values in the mid-19th century.

## 7 Conclusion and recommendations

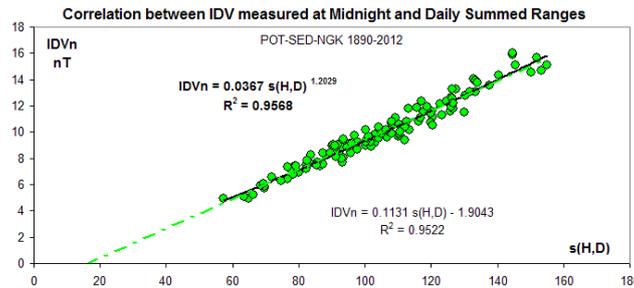
Using several lines of evidence we have shown that the scale values for the Helsinki magnetic data are in error at times. For  $H$ , the scale value for the interval 1866–1874 is too low by  $\sim 30\%$  and for  $D$  too low during 1886 by a factor of  $\sim 2$ . Lockwood (2013) reminds us about “the great importance of knowing, as far as is possible, the true provenance of historic data and of all the corrections and changes that may have subsequently been applied to them”. This is of critical importance in this situation where there is but one station for several years with usable data, so we urge Nevanlinna and colleagues to continue to re-examine the original data and their reduction. We urged M. Lockwood et al. (personal communication, 2013) to revise accordingly their analysis and derivation of IDV look-alikes where based on the Helsinki data. Such correction and revision has now happened: in Lockwood et al. (2014b) they write “a correction to seven years’ IDV(1d) data during solar cycle 11 was discussed, checked against newly-analysed independent data from St Petersburg and implemented”. This shows the value and power of comparisons with other data by the authors themselves and independently by other researchers.

Appendix A

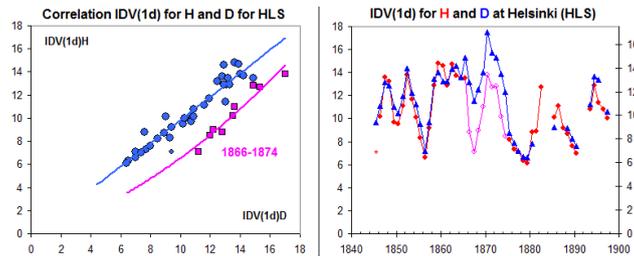
In this section we collect various Figures providing supplementary support for the analysis in the paper.



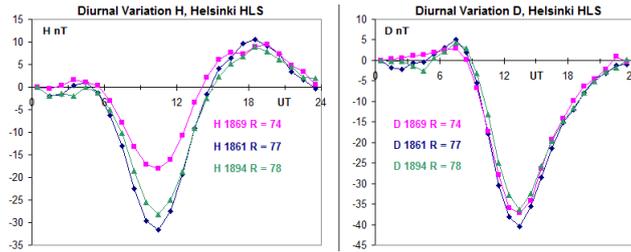
**Figure A1.** The average  $s(H)$  for each year is plotted against the average  $s(D)$  for that year. The data can be fitted to a power law as shown, which “explains” 96 % of the correlation. We use power laws because most regression plots show somewhat curved point clouds (“rivers” is probably a more descriptive term).



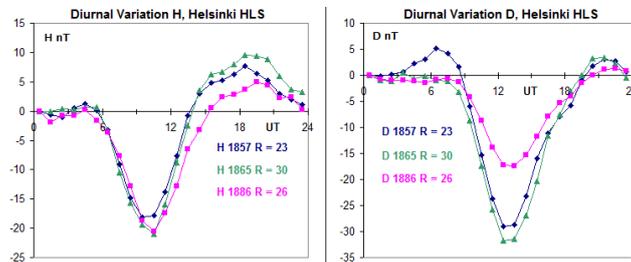
**Figure A2.** Correlation between yearly values of IDVn (midnight) and the average summed ranges for the day for  $H$  and  $D$ ,  $s(H, D)$ , for the POT-SED-NGK composite series for the interval 1890–2012. The dashed line is the linear relation extrapolated to vanishing IDVn.



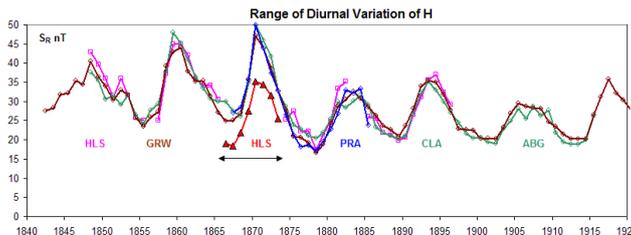
**Figure A3.** Using  $s(D)$  and  $s(H)$  we can calculate the corresponding IDV(1d) values for HLS (right-hand panel). There is generally a good agreement between values derived derived using  $s(H)$  and  $s(D)$ , except for the interval 1866–1874 for  $s(H)$  as also shown in the left-hand panel, plotting IDV(1d) derived from  $s(H)$  against IDV(1d) derived from  $s(D)$ . Pink symbols are for 1866–1874 while blue symbols are for data outside of that interval.



**Figure A4.** Diurnal variation of  $H$  (left) and of  $D$  (right) at Helsinki for 3 years with a SIDC sunspot number of  $\sim 75$ . It is very hard to escape the conclusion that the range of  $H$  for the year 1869 is too small compared to the other years with a similar sunspot number as the range of  $D$  is about the same for the 3 years shown. The timing from 1882 on has been adjusted to 1 h later (also in the following figure) because of a change from Göttingen time to local Helsinki time, which is not reflected correctly in the published data.



**Figure A5.** Diurnal variation of  $H$  (left) and of  $D$  (right) at Helsinki for 3 years with a SIDC sunspot number of  $\sim 25$ . It is very hard to escape the conclusion that the range of  $D$  for the year 1886 is too small compared to the other years with a similar sunspot number as the range of  $H$  is about the same for the 3 years shown. Detailed analysis shows that the problem exists from October 1885 to May 1887. In August 2003 we emailed Nevanlinna alerting him to this problem, but, unfortunately, no corrective action resulted from this. It is now clear that scale-value problems exist for both the  $H$  component and for the declination and that corrective action is mandatory before use of the Helsinki data.



**Figure A6.** The diurnal range (in nT) of the horizontal force for Prague (PRA) blue, Colaba (CLA)+Alibag (ABG) green, Greenwich (GRW) brown, and Helsinki (HLS) pink. For the interval 1866–1874, HLS (red triangles) is clearly seriously too low.

*Acknowledgements.* We acknowledge the impetus from participating in work of the ISSI International Team 233 and from serving as reviewers of Lockwood et al. (2013a, b). The Helsinki data was downloaded from <http://wdc.kugi.kyoto-u.ac.jp/> and cleaned for the usual types of errors found in such data sets, like data for 31 November, time being off by 1 h from 1882 on, and base values for declination being expressed in tens of minutes rather than the (correct) degrees. L. Svalgaard appreciates the continuing support from Stanford University.

Topical Editor C. Owen thanks T. L. Hansen and two anonymous referees for their help in evaluating this paper.

## References

- Bartels, J.: Archiv des Erdmagnetismus, Heft 5, Abh. Met. Inst., 8, Berlin, 1925.
- Bartels, J.: Terrestrial-magnetic activity and its relations to solar phenomena, *Terr. Magn. Atmos. Elec.*, 37, 1–52, 1932.
- Bartels, J.: Geomagnetic data on variations of solar radiation, 1, Wave-radiation, *Terr. Magn. Atmos. Elec.*, 51, 181–242, 1946.
- Emery, B. A., Coumans, V., Evans, D. S., Germany, G. A., Greer, M. S., Holeman, E., Kadinsky-Cade, K., Rich, F. J., and Xu, W.: Seasonal, Kp, solar wind, and solar flux variations in long-term single-pass satellite estimates of electron and ion auroral hemispheric power, *J. Geophys. Res.*, 113, A06311, doi:10.1029/2007JA012866, 2008.
- Lockwood, M.: Reconstruction and Prediction of Variations in the Open Solar Magnetic Flux and Interplanetary Conditions, *Living Rev. Solar Phys.* 10, 4, doi:10.12942/lrsp-2013-4, 2013.
- Lockwood, M. and Owens, M. J.: Centennial changes in the heliospheric magnetic field and open solar flux: The consensus view from geomagnetic data and cosmogenic isotopes and its implications, *J. Geophys. Res.*, 116, A04109, doi:10.1029/2010JA016220, 2011.
- Lockwood, M. and Owens, M. J.: Implications of the Recent Low Solar Minimum for the Solar Wind During the Maunder Minimum, *Astrophys. J. Lett.*, 781, 1, doi:10.1088/2041-8205/781/1/L7, 2014.
- Lockwood, M., Stamper, R., and Wild, M. N.: A doubling of the Sun's coronal magnetic field during the past 100 years, *Nature*, 399, 437–439, doi:10.1038/20867, 1999.
- Lockwood, M., Barnard, L., Nevanlinna, H., Owens, M. J., Harrison, R. G., Rouillard, A. P., and Davis, C. J.: Reconstruction of geomagnetic activity and near-Earth interplanetary conditions over the past 167 yr – Part 1: A new geomagnetic data composite, *Ann. Geophys.*, 31, 1957–1977, doi:10.5194/angeo-31-1957-2013, 2013a.
- Lockwood, M., Barnard, L., Nevanlinna, H., Owens, M. J., Harrison, R. G., Rouillard, A. P., and Davis, C. J.: Reconstruction of geomagnetic activity and near-Earth interplanetary conditions over the past 167 yr – Part 2: A new reconstruction of the interplanetary magnetic field, *Ann. Geophys.*, 31, 1979–1992, doi:10.5194/angeo-31-1979-2013, 2013b.
- Lockwood, M., Nevanlinna, H., Vokhmyanin, M., Ponyavin, D., Sokolov, S., Barnard, L., Owens, M. J., Harrison, R. G., Rouillard, A. P., and Scott, C. J.: Reconstruction of geomagnetic activity and near-Earth interplanetary conditions over the past 167 yr – Part 3: Improved representation of solar cycle 11, *Ann. Geophys.*, 32, 367–381, doi:10.5194/angeo-32-367-2014, 2014a.
- Lockwood, M., Nevanlinna, H., Barnard, L., Owens, M. J., Harrison, R. G., Rouillard, A. P., and Scott, C. J.: Reconstruction of geomagnetic activity and near-Earth interplanetary conditions over the past 167 yr – Part 4: Near-Earth solar wind speed, IMF, and open solar flux, *Ann. Geophys.*, 32, 383–399, doi:10.5194/angeo-32-383-2014, 2014b.
- Mayaud, P. N.: Derivation, Meaning, and Use of Geomagnetic Indices, *Geophys. Monogr. Ser.*, Vol. 22, 154 pp., ISBN 0-87590-022-4, American Geophysical Union, Washington D.C., 1980.
- Morozova, A. L., Ribeiro, P., and Pais, M. A.: Correction of artificial jumps in the historical geomagnetic measurements of Coimbra Observatory, Portugal, *Ann. Geophys.*, 32, 19–40, doi:10.5194/angeo-32-19-2014, 2014.
- Moos, N. A. F.: Colaba Magnetic Data, 1846 to 1905, 2, The Phenomenon and its Discussion, 782 pp., Central Govt. Press, Bombay, 1910.
- Nervander, J. J.: Observations faites à l'observatoire Magnétique et Météorologique de Helsingfors 1844–1848, Vol. I–IV, Helsingfors, 1850.
- Nevanlinna, H.: Results of the Helsinki magnetic observatory 1844–1912, *Ann. Geophys.*, 22, 1691–1704, doi:10.5194/angeo-22-1691-2004, 2004.
- Nevanlinna, H., Ketola, A., and Kangas, T.: Magnetic Results from Helsinki Magnetic-Meteorological Observatory, *Finn. Met. Inst., Geophys. Publ.*, 27, 155 pp., 1992.
- Svalgaard, L.: Building a Sunspot Group Number Backbone Series, available at: <http://www.leif.org/research/SSN/Svalgaard11.pdf> (last access: 3 June 2014), 2013.
- Svalgaard, L. and Cliver, E. W.: New Geomagnetic Index (IDV) Measuring Magnitude of Interplanetary Magnetic Field, AGU Fall SH21B-0108, available at: <http://www.leif.org/research/AGUFall2003SH21B-0108.pdf> (last access: 3 June 2014), 2003.
- Svalgaard, L. and Cliver, E. W.: The IDV index: Its derivation and use in inferring long-term variations of the interplanetary magnetic field strength, *J. Geophys. Res.*, 110, A12103, doi:10.1029/2005JA011203, 2005.
- Svalgaard, L. and Cliver, E. W.: Interhourly variability index of geomagnetic activity and its use in deriving the long-term variation of solar wind speed, *J. Geophys. Res.*, 112, A10111, doi:10.1029/2007JA012437, 2007.
- Svalgaard, L. and Cliver, E. W.: Heliospheric magnetic field 1835–2009, *J. Geophys. Res.*, 115, A09111, doi:10.1029/2009JA015069, 2010.
- Svalgaard, L. and Cliver, E. W.: Update: Heliospheric magnetic field 1835–2013, *J. Geophys. Res.*, in preparation, 2014.
- Svalgaard, L., Cliver, E. W., and Le Sager, P.: Determination of interplanetary magnetic field strength, solar wind speed, and EUV irradiance, 1890–2003, in *Proceedings of ISCS 2003 Symposium: Solar Variability as an Input to the Earth's Environment*, Eur. Space Agency Spec. Publ., ESA SP-535, 15, 2003.
- Svalgaard, L., Cliver, E. W., and Le Sager, P.: IHV: A new long-term geomagnetic index, *Adv. Space. Res.*, 34, 1112–1120, doi:10.1016/j.asr.2007.06.066, 2004.
- Wolf, J. R.: Neue Untersuchungen über die Periode der Sonnenflecken und ihre Bedeutung, *Viertel. Natur. Ges. Bern*, 245, 249–270, Bern, 1852.