

Comments on

An analysis of heliospheric magnetic field flux based on sunspot number from 1749 to today and prediction for the coming solar minimum

Molly L. Goelzer,^{1,2} Charles W. Smith,¹ Nathan A. Schwadron,¹ and K. G. McCracken³

JOURNAL OF GEOPHYSICAL RESEARCH: SPACE PHYSICS, VOL. 118, 7525–7531, doi:10.1002/2013JA019404, 2013

Leif Svalgaard

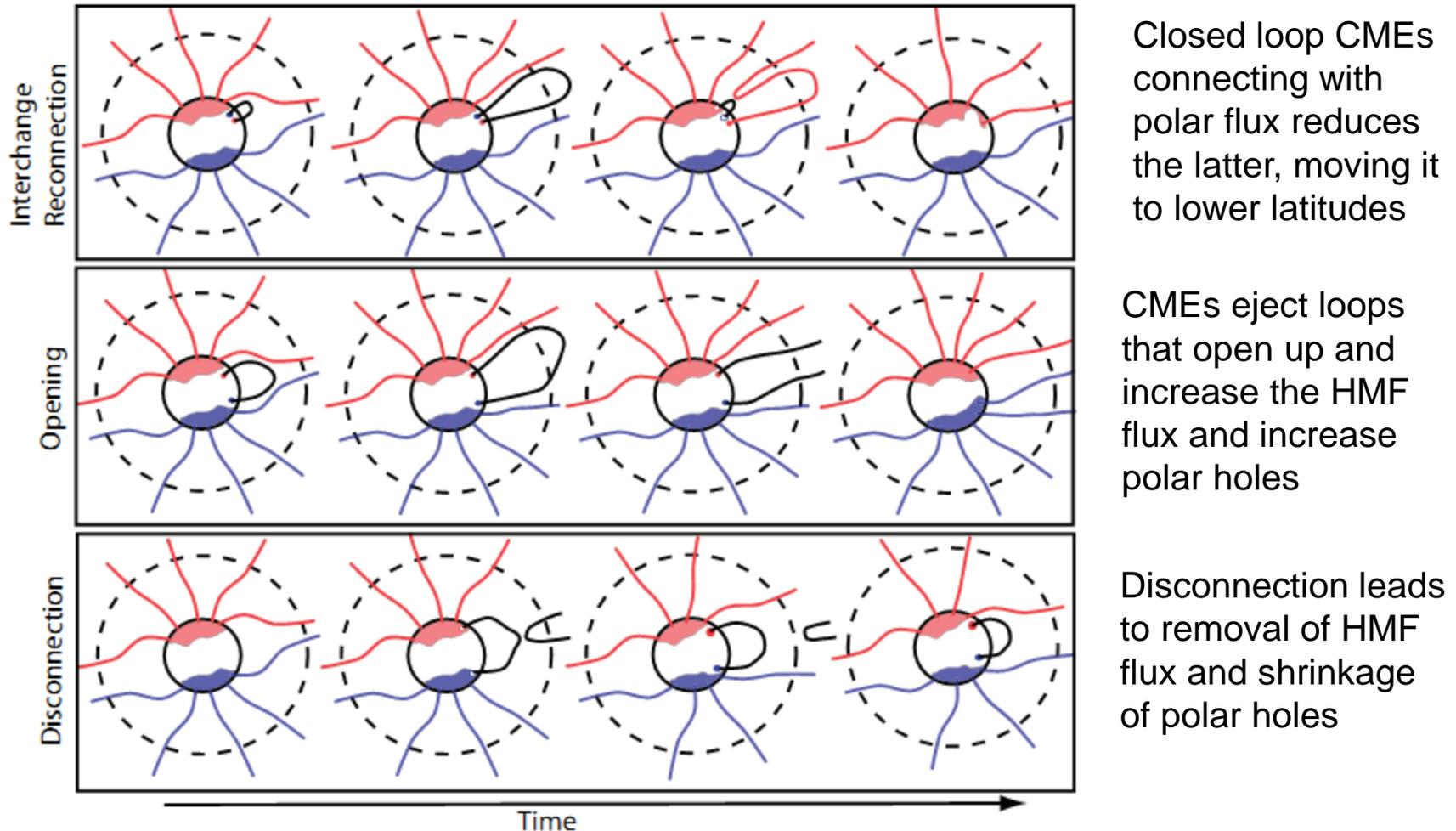
Journal Club Discussion

Stanford University

Wed. 19 Feb. 2014

Magnetic Flux Balance in the Heliosphere

Schwadron et al. ApJ 722, L132, 2010



Closed loop CMEs connecting with polar flux reduces the latter, moving it to lower latitudes

CMEs eject loops that open up and increase the HMF flux and increase polar holes

Disconnection leads to removal of HMF flux and shrinkage of polar holes

The CME Flux Rate

The theory posits two components of the HMF: the CME associated magnetic flux ϕ_{CME} from the ejecta and the open magnetic flux Φ_0 of the steady solar wind. The time derivative of the CME-associated flux Φ_{ej} is written as

$$\frac{d\Phi_{\text{ej}}}{dt} = f(1 - D)\phi_{\text{CME}} - \Phi_{\text{ej}} \left(\frac{1}{\tau_{ic}} + \frac{1}{\tau_d} + \frac{1}{\tau_o} \right)$$

A large fraction, $D \approx 1/2$, of this flux opens as the CME launches from the Sun, so the closed flux created by an individual CME is $(1 - D)\phi_{\text{CME}}$. The other factor that enters the source term for closed ejecta-associated flux is the frequency, $f(t)$, of CME ejection, which depends on time, t , varying typically from 0.5 day^{-1} near solar minimum to 3 day^{-1} near solar maximum.

The loss term has three parts: A decay of Φ_{ej} due to interchange reconnection on a characteristic timescale of τ_{ic} ; disconnection on a timescale of τ_d ; and opening of the flux loops when they merge into the background on a timescale of τ_o . All of these timescales are poorly known, but one can make guesses: $\tau_{ic} = 20\text{d}$, $\tau_d = 6\text{y}$, $\tau_o = 2.5\text{y}$.

The Open Flux Rate

For the magnetic flux not associated with CMEs [the 'open' flux Φ_0] they set the time derivative equal to the loss by disconnection (on timescale τ_d) subtracted from the source due to the opening of CME-associated flux (on timescale τ_o), thus:

$$\frac{d\Phi_0}{dt} = -\frac{\Phi_0 - \Phi_{\text{flr}}}{\tau_d} + \frac{\Phi_{\text{ej}}}{\tau_o}$$

If there is a 'floor' on which the CME-associated flux 'rides' then the disconnection loss-term should be measured relative to the floor Φ_{flr}

Adding the rate of change of the CME-associated flux, given by

$$\frac{d\Phi_{\text{ej}}}{dt} = f(1 - D)\phi_{\text{CME}} - \Phi_{\text{ej}} \left(\frac{1}{\tau_{ic}} + \frac{1}{\tau_d} + \frac{1}{\tau_o} \right)$$

And using $\Phi_{\text{tot}} = \Phi_0 + \Phi_{\text{ej}} + \Phi_{\text{flr}}$

The Total Flux Rate

For the magnetic flux not associated with CMEs [the 'open' flux Φ_0] they set the time derivative equal to the loss by disconnection (on timescale τ_d) subtracted from the source due to the opening of CME-associated flux (on timescale τ_o), thus:

$$\frac{d\Phi_0}{dt} = -\frac{\Phi_0 - \Phi_{\text{flr}}}{\tau_d} + \cancel{\frac{\Phi_{\text{ej}}}{\tau_o}}$$

If there is a 'floor' on which the CME-associated flux 'rides' then the disconnection loss-term should be measured relative to the floor Φ_{flr}

Adding the rate of change of the CME-associated flux, given by

$$\frac{d\Phi_{\text{ej}}}{dt} = f(1 - D)\phi_{\text{CME}} - \Phi_{\text{ej}} \left(\frac{1}{\tau_{\text{ic}}} + \frac{1}{\tau_d} + \cancel{\frac{1}{\tau_o}} \right)$$

The last members of each term cancel and we get the rate of change of the total flux:

$$\frac{d\Phi_{\text{tot}}}{dt} = -\frac{\Phi_{\text{tot}} - \Phi_{\text{flr}}}{\tau_d} + f(1 - D)\phi_{\text{CME}} - \frac{\Phi_{\text{ej}}}{\tau_{\text{ic}}}$$

Determining Total Hemispheric Flux

The integral solution for the ejecta-associated [CME] magnetic flux is

$$\Phi_{ej} = \int_{-\infty}^t dt' \exp\left(-\frac{t-t'}{\tau_c}\right) f(t')(1-D)\phi_{CME}$$

Where the characteristic loss-time of the closed [CME] flux is

$$\frac{1}{\tau_c} = \frac{1}{\tau_{ic}} + \frac{1}{\tau_o} + \frac{1}{\tau_d} = 1/(19.4 \text{ days}) = 1/(4.3 \text{ AU-time})$$

And where the CME rate $f(t)$ is derived from the Sunspot Number SSN:

$$f(t) = SSN(t) / 25$$

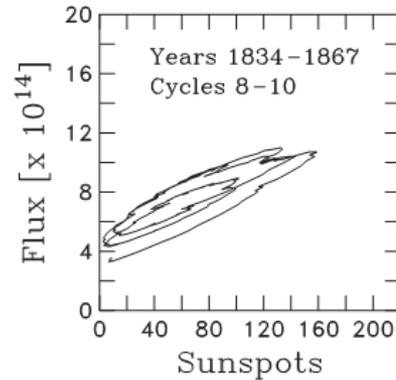
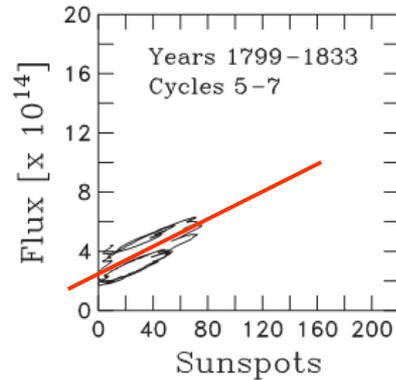
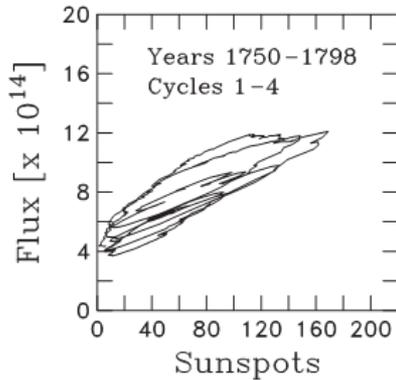
The integral solution for 'open' heliospheric magnetic flux is

$$\Phi_0 = \Phi_{fir} + \int_{-\infty}^t dt' \exp\left(-\frac{t-t'}{\tau_d}\right) \frac{\Phi_{ej}}{\tau_o}$$

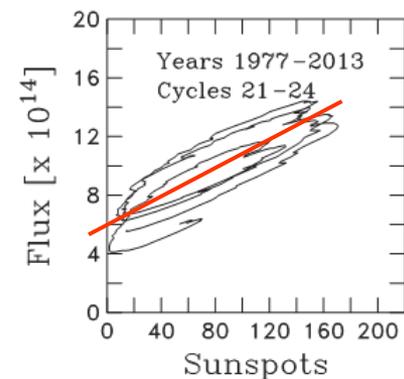
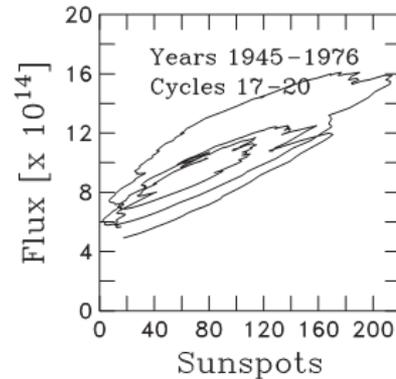
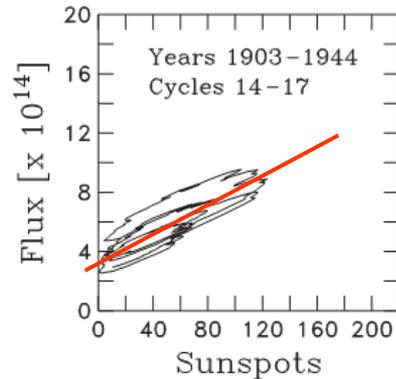
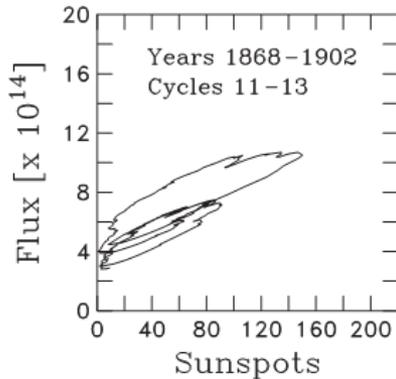
The total flux becomes $\Phi_{tot} = \Phi_0 + \Phi_{ej} = \oint \mathbf{B}_P \cdot \hat{n} dS = 4\pi R^2 |B_P|$

Which evaluated for $R = 1 \text{ AU}$ allows you to infer the HMF field strength, B , at Earth. The subscript P in B_P stands for the 'Parker Spiral Field'.

Calculation of Total Flux from SSN

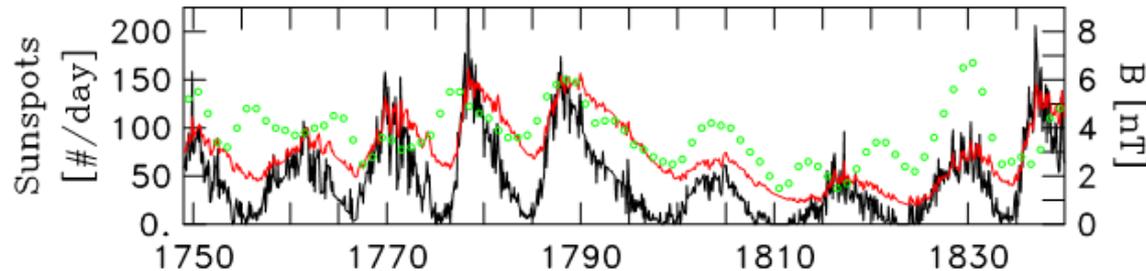


Plot of total flux as a function of the SSN, for several groups of cycles. Note the different 'floors'.

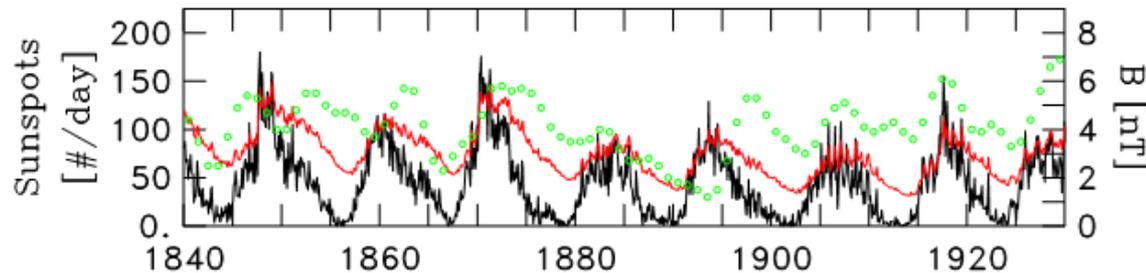


Because the flux rises quickly with sunspots, which they use as a proxy for CME activity, and falls more slowly as sunspot activity decreases, there is a noted hysteresis effect. In the calculation the 'floor' Φ_{flr} was set to zero, contrary to observations.

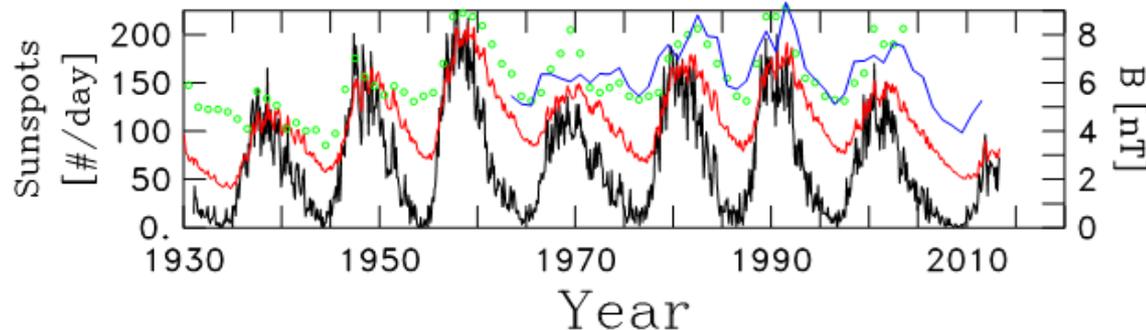
Comparing Theory with Observations



Black is Official Sunspot Number SSN from SIDC
Red is B_p calculated from their theory



Green is B deduced from ^{10}Be data by McCracken 2007

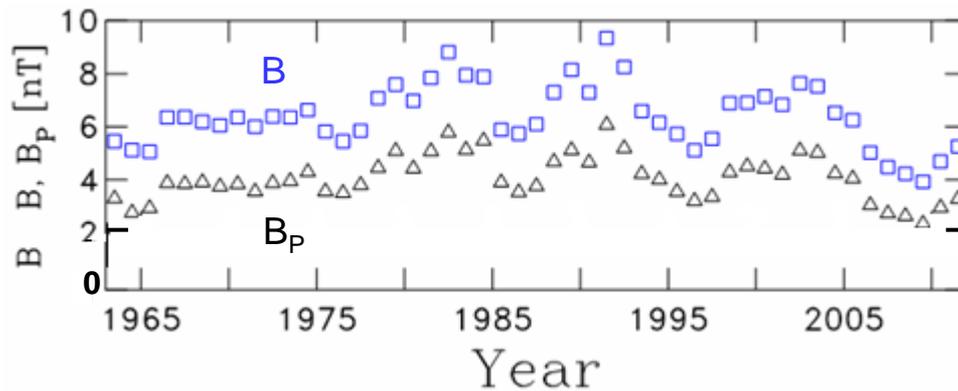


Blue is B taken from the spacecraft-based OMNI dataset

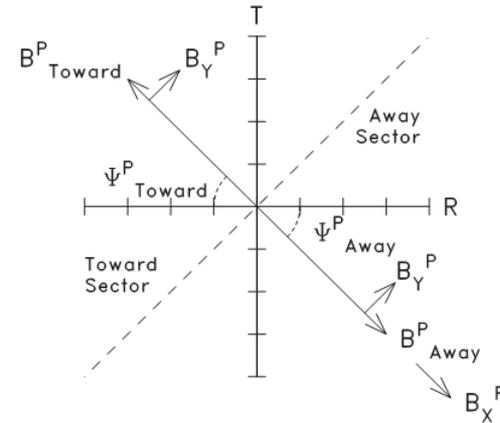
At first blush the correspondences don't look too good...

'Explanation' in terms of Difference between the Ideal 'Parker Field' and the Actual, Messy, Observed Field

Goelzer et al. state that the theory predicts the 'Parker field' $|B_P|$ while the values for $|B|$ also contain the azimuthal fields often associated with magnetic clouds and the turbulent magnetic fluctuations, both of which are absent from the definition of $|B_P|$. Therefore, they fully expect that $|B_P| < |B|$



THE ASTROPHYSICAL JOURNAL, 727:8 (6pp), 2011 January 20
 © 2011. The American Astronomical Society. All rights reserved. Printed in the U.S.A.



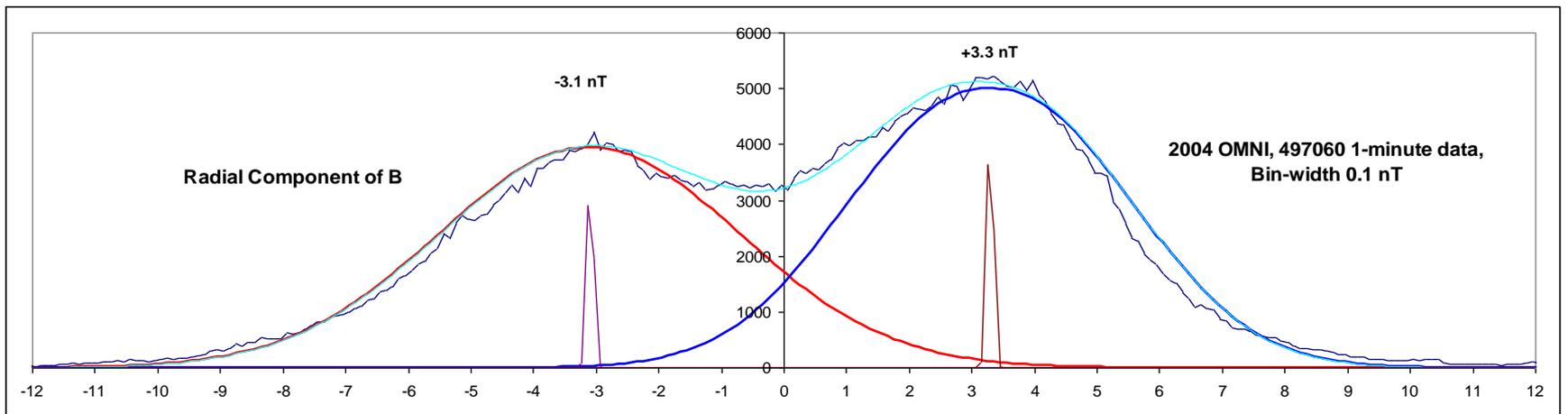
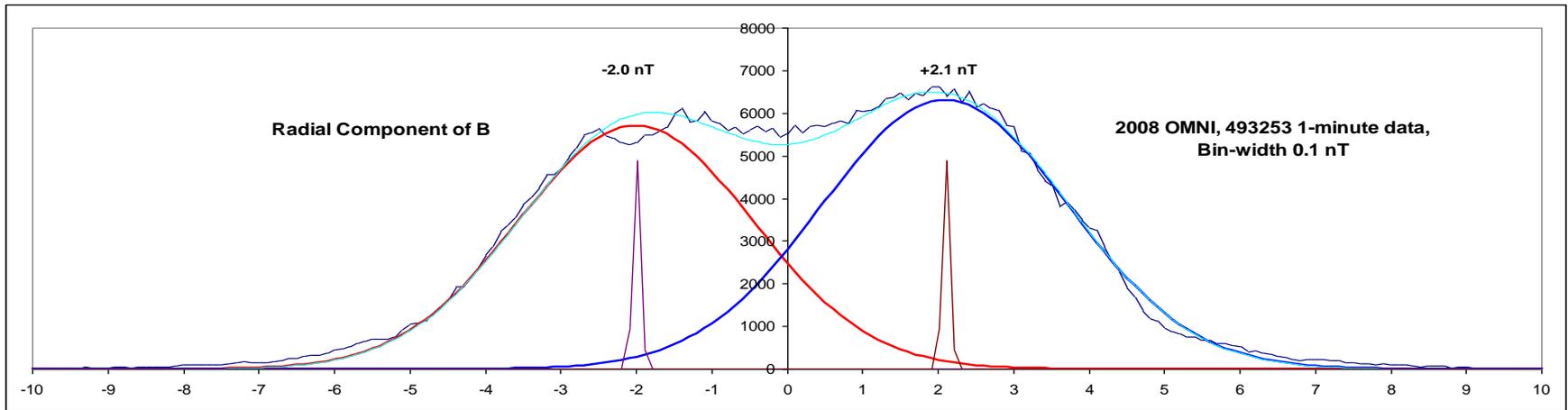
doi:10.1088/0004-637X/727/1/8

INTERPLANETARY MAGNETIC FLUX DEPLETION DURING PROTRACTED SOLAR MINIMA

DAVID E. CONNICK, CHARLES W. SMITH, AND NATHAN A. SCHWADRON

Attempts to compute B_P from OMNI data and then compare with B

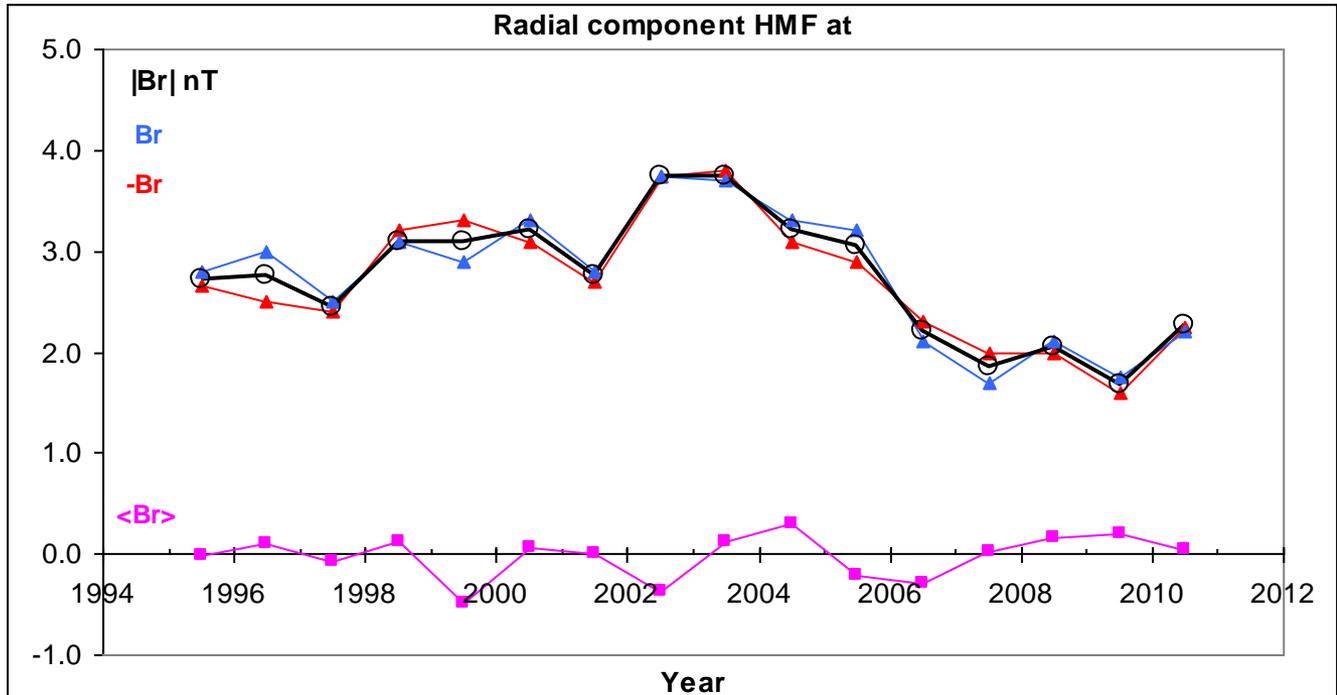
Simpler method: Looking at 1-minute values of the radial component, B_r , it appears that they can be described by two overlapping Gaussians, e.g.:



Fitting two Gaussians to the distributions gives us the most probable value of B_r . We can do this for each year since 1995 for which we have 1-min values

Yearly values of |Br|

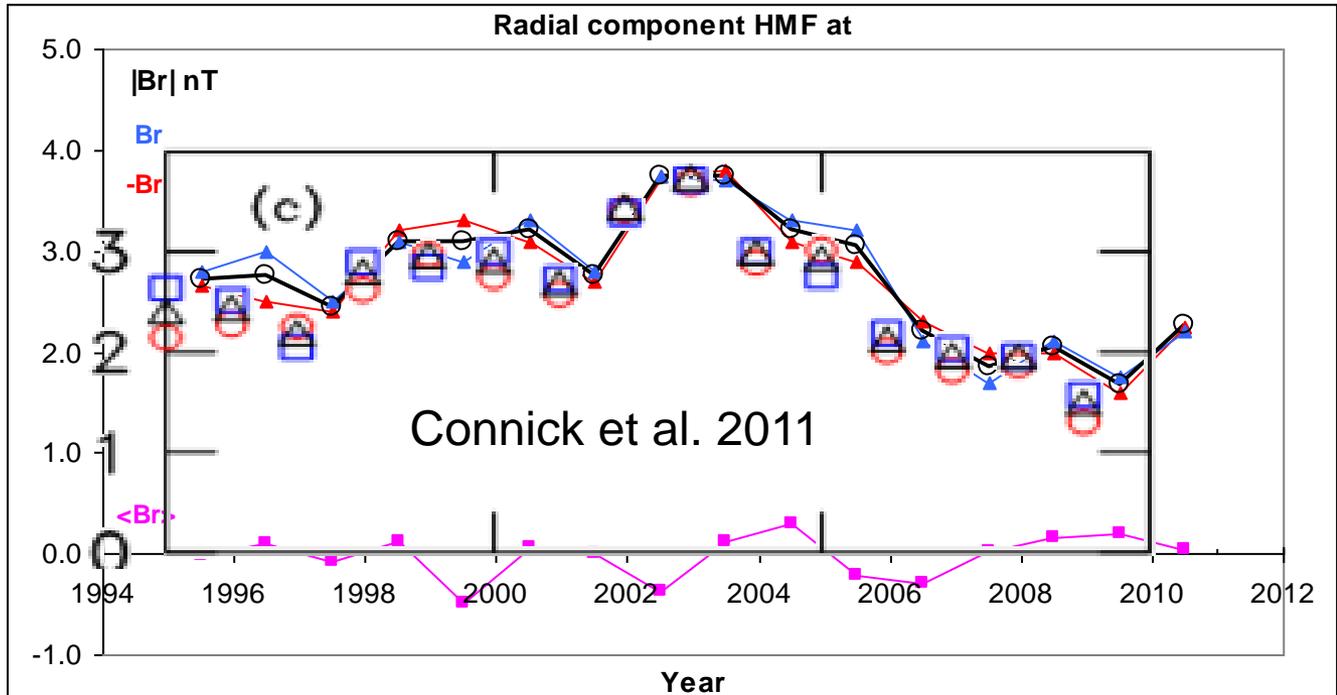
| Year | Br |
|--------|-------|
| 1995.5 | 2.725 |
| 1996.5 | 2.75 |
| 1997.5 | 2.45 |
| 1998.5 | 3.1 |
| 1999.5 | 3.1 |
| 2000.5 | 3.2 |
| 2001.5 | 2.75 |
| 2002.5 | 3.75 |
| 2003.5 | 3.75 |
| 2004.5 | 3.2 |
| 2005.5 | 3.05 |
| 2006.5 | 2.2 |
| 2007.5 | 1.85 |
| 2008.5 | 2.05 |
| 2009.5 | 1.675 |
| 2010.5 | 2.275 |



Blue symbols: +Br; Red: -Br; Pink: overall signed average for year

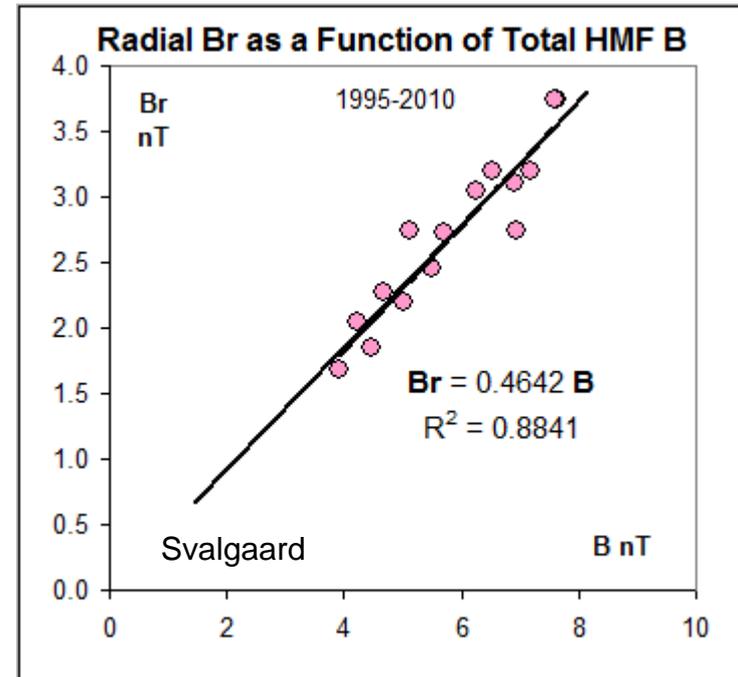
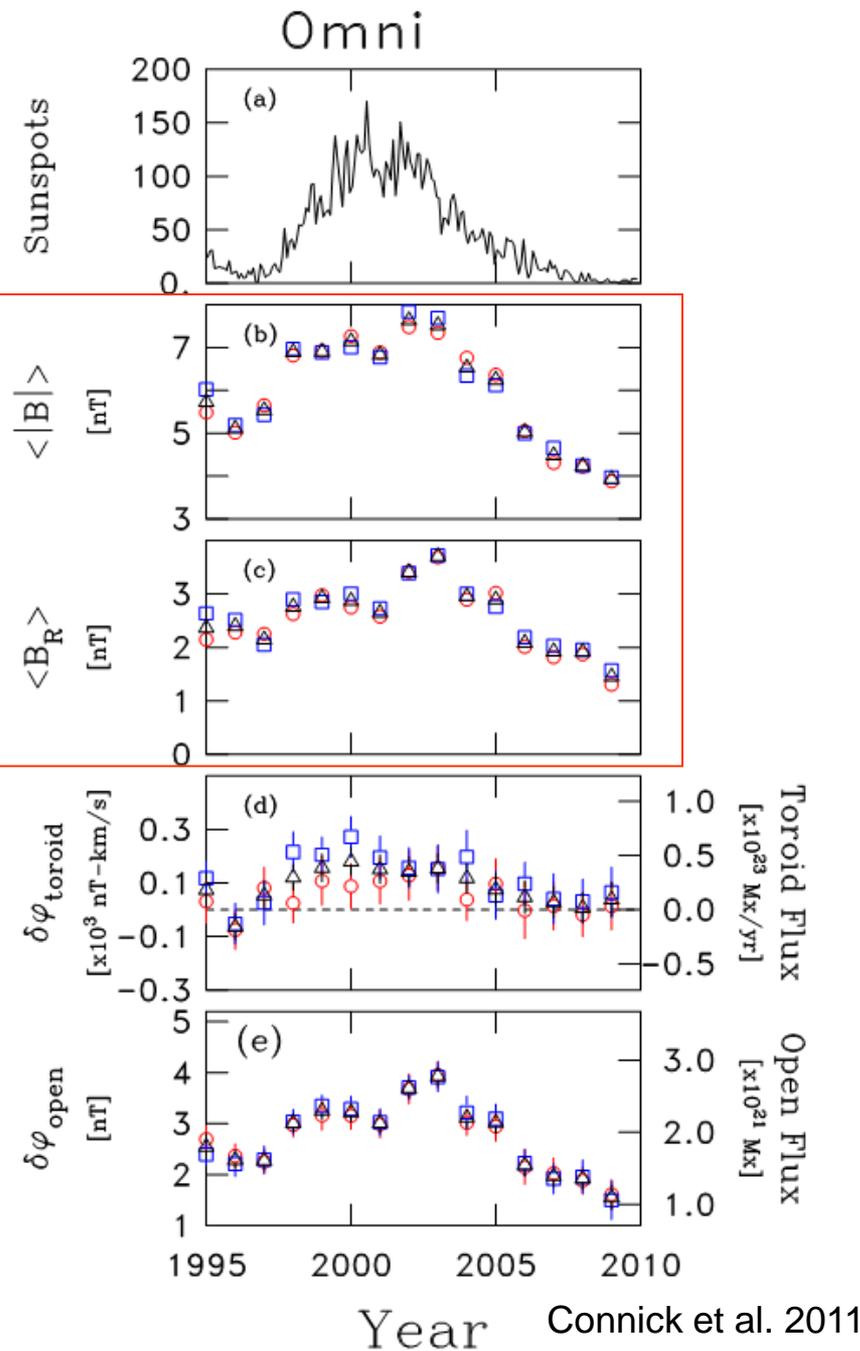
Yearly values of |Br|

| Year | Br |
|--------|-------|
| 1995.5 | 2.725 |
| 1996.5 | 2.75 |
| 1997.5 | 2.45 |
| 1998.5 | 3.1 |
| 1999.5 | 3.1 |
| 2000.5 | 3.2 |
| 2001.5 | 2.75 |
| 2002.5 | 3.75 |
| 2003.5 | 3.75 |
| 2004.5 | 3.2 |
| 2005.5 | 3.05 |
| 2006.5 | 2.2 |
| 2007.5 | 1.85 |
| 2008.5 | 2.05 |
| 2009.5 | 1.675 |
| 2010.5 | 2.275 |



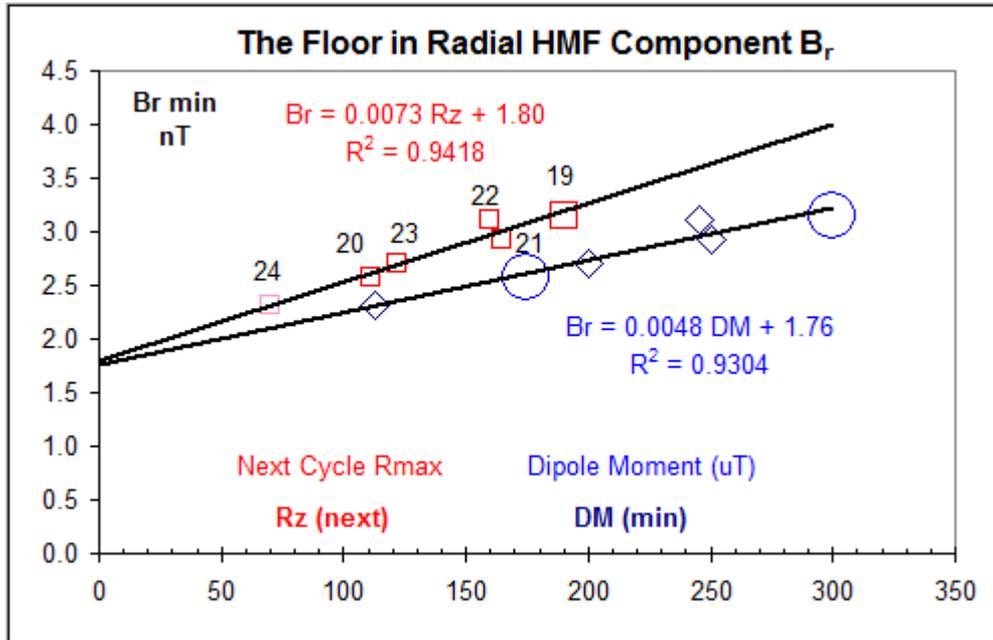
Blue symbols: +Br; Red: -Br; Pink: overall signed average for year

Estimation of B



Connick et al. and I agree what the observed values of B_r and B are in spite of the difference in methodology

Some (of my) Prediction Numerology



$$B_r = 0.46 B_{min}$$

$$B_r = 1.80 + 0.0073 R_{next}$$

$$0.46 B_{min} = 1.80 + 0.0073 R_{next}$$

$$B_{min} = 3.87 + 0.0158 R_{next}$$

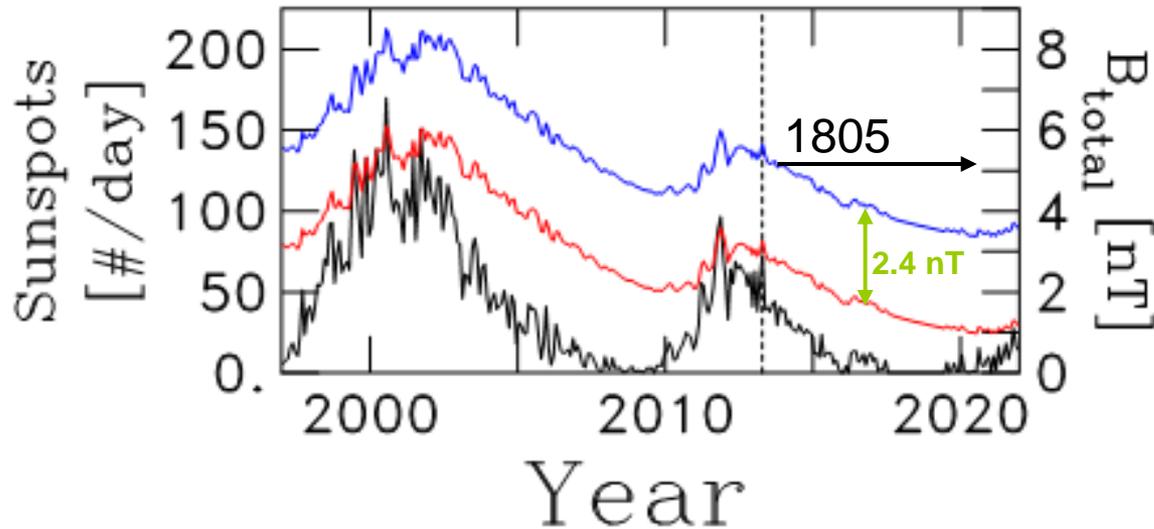
Hence prediction of R_{next} :

$$R_{next} = (B_{min} - 3.87)/0.0158$$

With the Floor at 3.87 nT

We plot the radial component B_r at minimum as a 'function' of the maximum sunspot number R_{next} for the next cycle on the assumption that the HMF at that time is a precursor for the sunspot cycle. B_r at minimum seems to depend on the 'Dipole Moment' of the global solar magnetic field (polar fields) which so far *has* been shown to a decent predictor of R_{next} . A problem is that we don't know B_{min} until the next minimum.

How to Guess the next B_{\min} ?



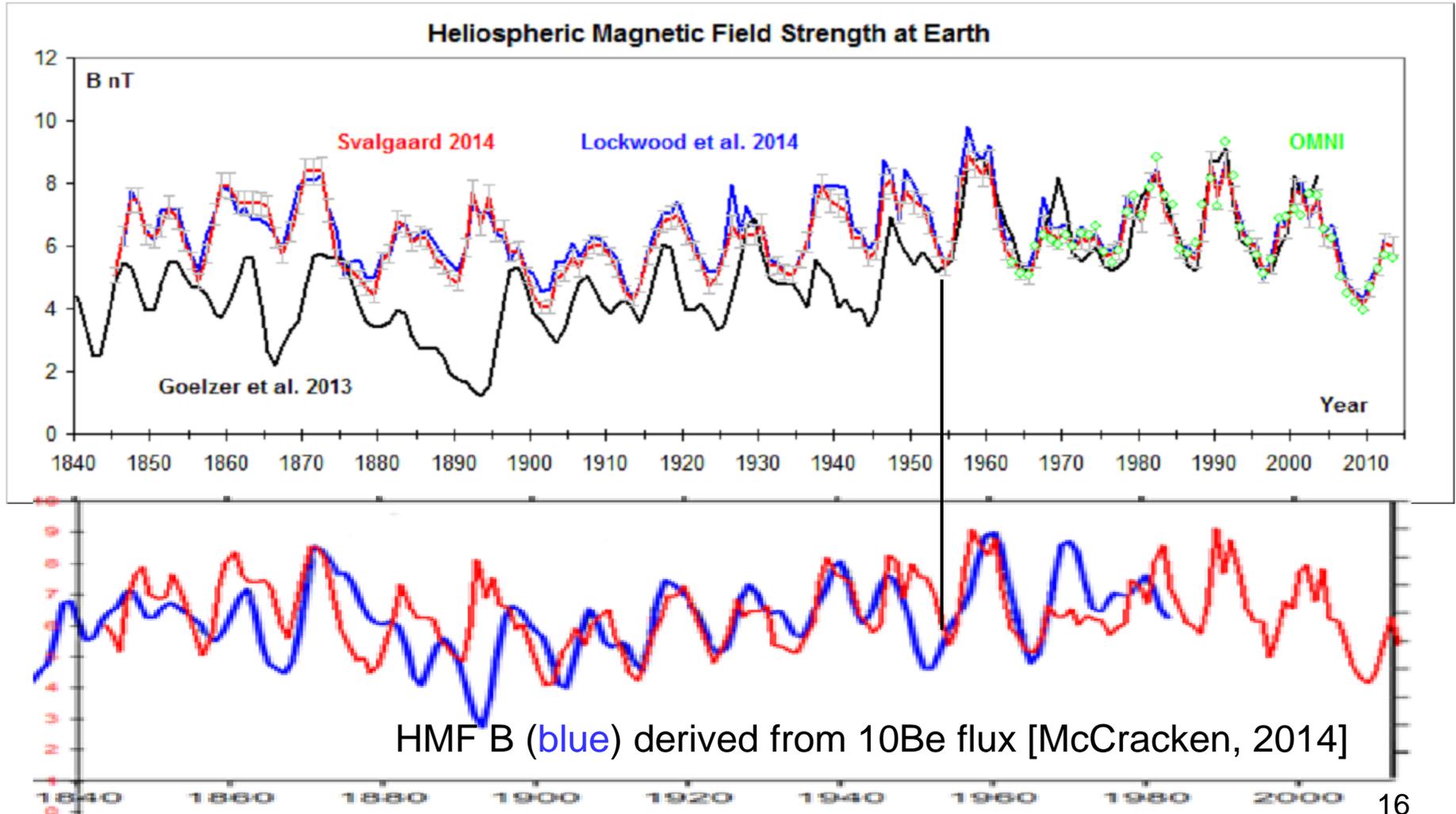
Goelzer et al. do so by noting that early in 2013 marks the peak in the solar cycle and that the sunspot number is comparable to what was seen during the Dalton Minimum. The years following 1805 thereby serve as a prediction for the coming 10 years of solar activity.

“From year 2013 onward the sunspot number is obtained from the historical record 1805 onward. The resulting $|\mathbf{B}_p|$ for 2020 shown in red is 1 nT lower than in the last protracted solar minimum. Prediction for $|\mathbf{B}|$ shown in green [?] is based on the observation that $|\mathbf{B}| - |\mathbf{B}_p|$ averages 2.4 nT, although it is less during solar minimum.”

All of this hinges on the SSNs going in and on the ^{10}Be HMF \mathbf{B} used to validate the theory are both correct.

And Therein Lies a Problem

Recent reconstructions (that we discussed last week) of HMF B are strongly at variance with the values used by Goelzer et al:



True Collaborative Spirit

I brought this problem to the attention of the authors with some suggestions [and data] to re-do the analysis with better input data and after some back-and-forth got this e-mail:

Nathan Schwadron <nschwadron@guero.sr.unh.edu>

Mon, Feb 17, 2014 at 10:20 AM To: Leif Svalgaard <leif@leif.org>

Cc: Charles Smith <charles.smith@unh.edu>, Michael Lockwood <m.lockwood@reading.ac.uk>, Ken McCracken <jellore@hinet.net.au>, Nathan Schwadron <n.schwadron@unh.edu>, Molly Goelzer <mlw292@wildcats.unh.edu>, Matthew James Owens <m.j.owens@reading.ac.uk>

I am totally on board with following the work up as you suggest

So, this line of research is alive and well and interesting results are bound to follow, as other groups are also attempting to model the HMF, especially with its behavior during Grand Minima [such as the Maunder Minimum]. Here is my view of the data and relationships we must try to understand to further that goal:

